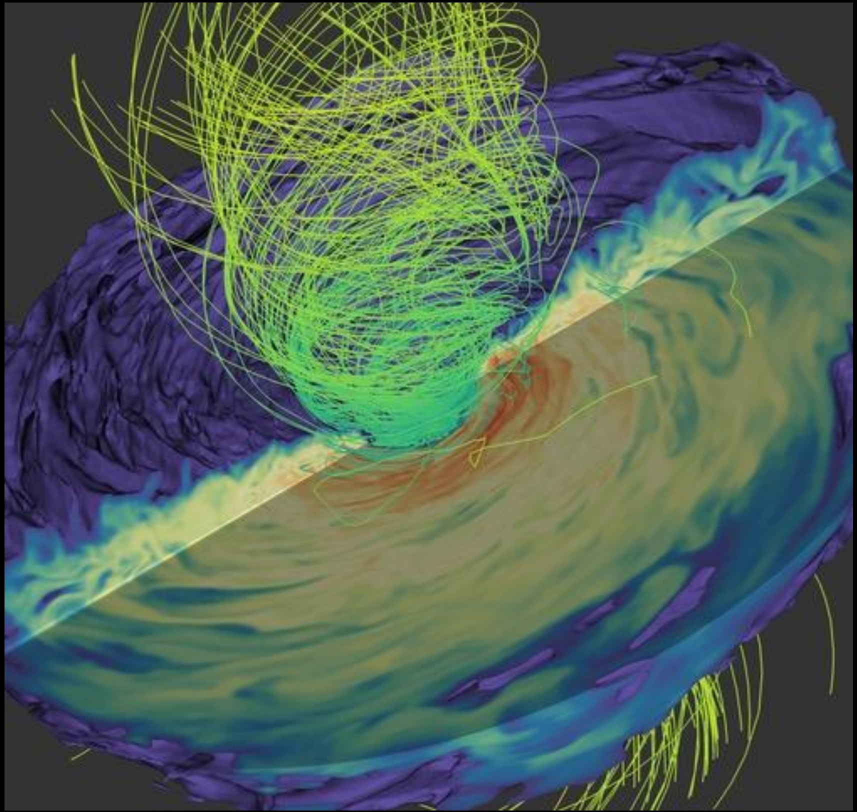
Simulations of Accreting Binary Supermassive Black Holes Approaching Merger



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KICC 16-20/09/2019
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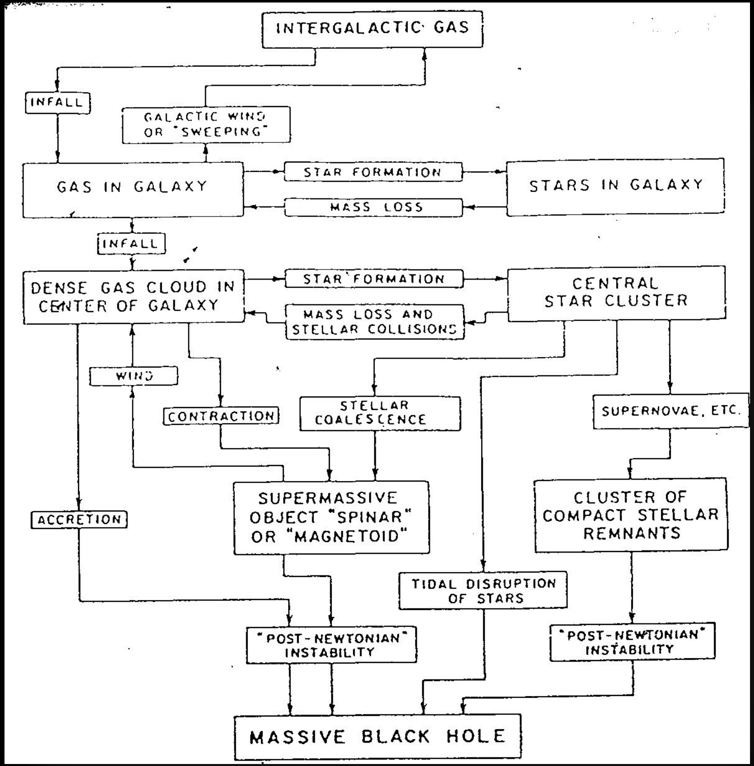
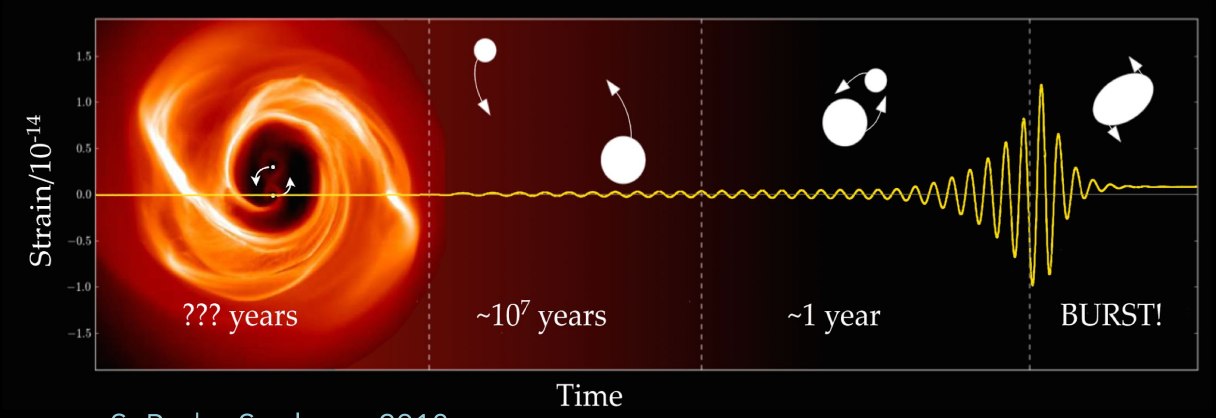
There are Black Holes at Centers of Galaxies!



Martin Rees, 1978

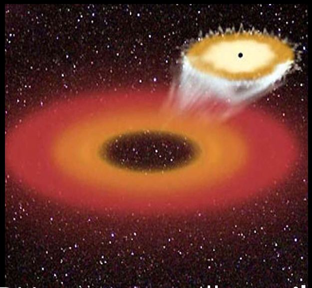
There might even be two …

What are the binary dynamics that take the BHs from galaxy merger scales to the GW scale? - Begelman, Blandford & Rees 1980



S. Burke-Spolaor +2018

- up to ~10% of the total mass is radiated in GW energy – e.g Campanelli+2006



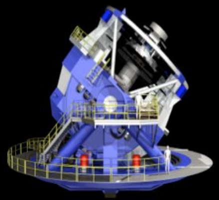
- The BH remnant will **recoil** from its host structure, depending on the BH spins and masses at merger – e.g Campanelli+2007 …

Assume that stellar dynamical friction, torques from gas, or other processes can bring the pair to sub-pc scales, then GW should do the rest …

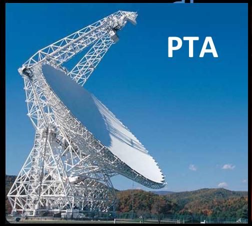
And they could be EM – bright …

- Supermassive BHs in AGN are surrounded by accreting hot gas and emit powerful radio jets!
- Binary supermassive BH are primary GW sources for **LISA** and **PTA** campaigns.
- As EM sources, they are ideal candidate for exploring plasma physics in the strongest and most dynamical regime of gravity.

Electromagne
tic



Gravitation
al



Holy Grail for
Multi-Messenger Astrophysics!

What are the EM signals associated with SMBBH merger?

Are they EM-Distinguishable from single AGN?

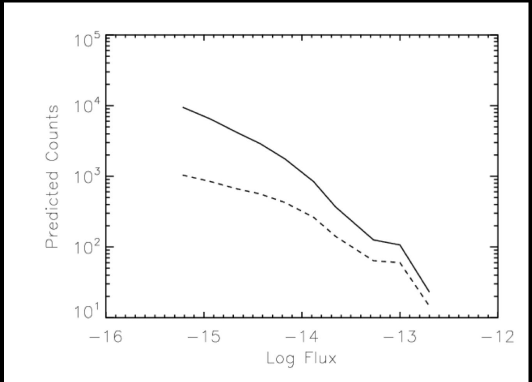
- There are two possible signals that can distinctively mark binaries ($15 < a/r_g < 10^3$) – Roedig, Krolik, Miller 2014:

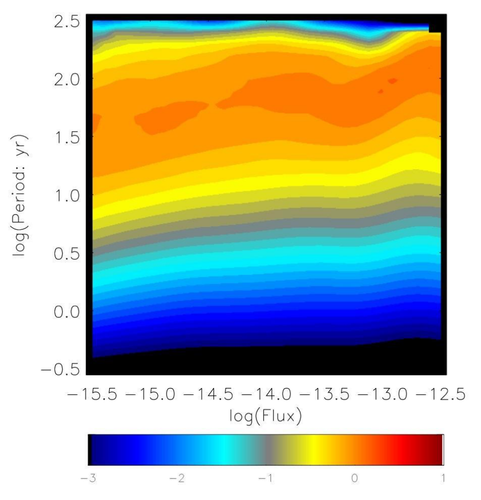
    - *Notch* in the IR/optical/UV spectrum
    - Periodically - modulated hard X-ray component

- Galaxy evolution simulations predict signal counts - Krolik, Volonteri, Dubois, and Devriendt, 2019

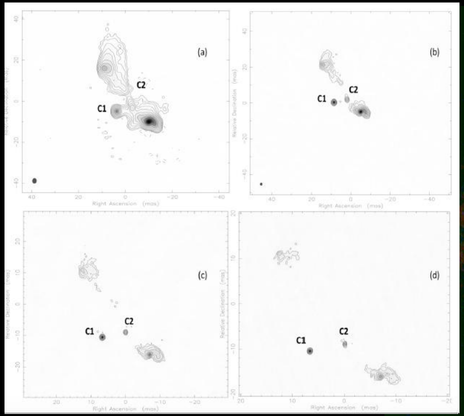
    - There might be $\sim 10^2$ with fluxes in the bands containing the signal $>10^{-13}$ erg cm^{-2} s^{-1}
    - The distribution peaks for orbital periods ~20--200 yr, but **~ 10% have periods ~ 3-5 yr**.

That is in the range that can be probed by PTA!

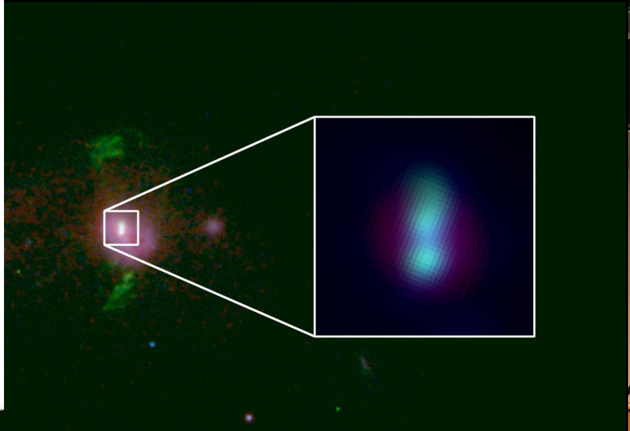




So far, a handful of Candidates …

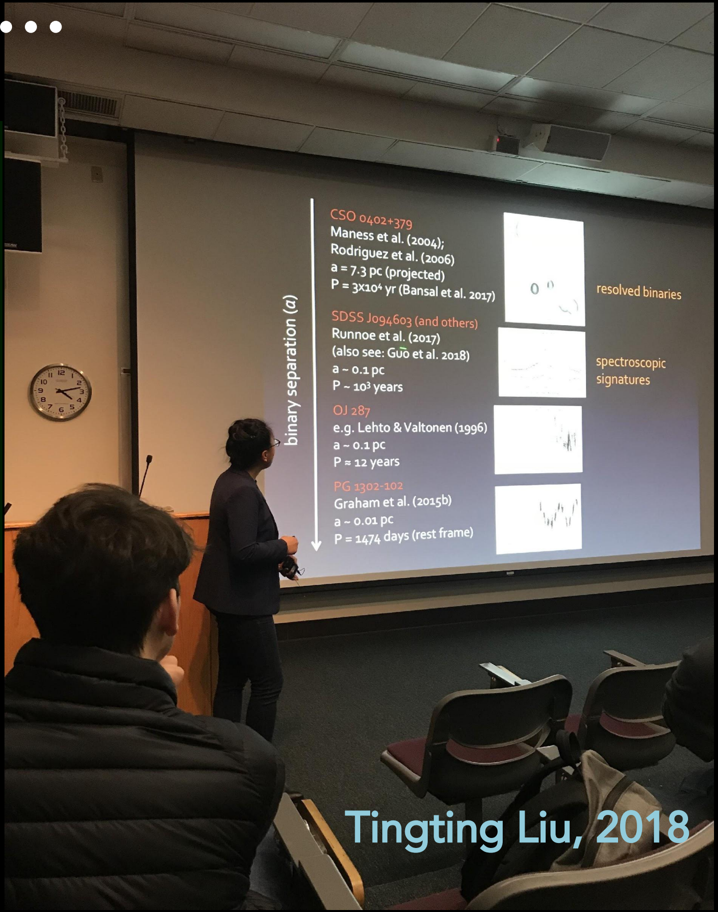


Radio galaxy 0402+379 - Bansal+2017, 12 years of multi-frequency VLBI observations



Goulding+ ApJL 2019;
HST image of SDSS J1010+1413
PTA source

SDSS discovered ~ 10^6 AGN; Future astronomical surveys, e.g. LSST will study optical variability in a larger sample, so "many" binary-AGN are expected to be uncovered in the haystack!



Tingting Liu, 2018

Modeling Merging Supermassive Black Hole Binaries

Realistic simulations of the last stages of the merger are needed for EM identification and characterization!

What this requires?

Choose astrophysically motivated disk models, use GR-MHD, "realistic" thermodynamics and radiation treatment, run for **long enough** to equilibrate the system while **resolving MRI** for proper angular momentum transport in the gas and close to the **BH horizons – all, considering that the spacetime is dynamically changing!**

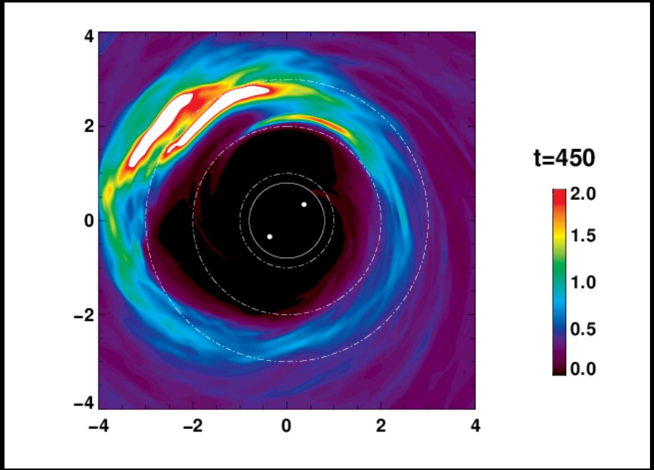
How much gas is present at Merger?

- Early Newtonian HD in 1D found little or no accretion close to the binary, as binary torques carve a nearly empty cavity of ~ 2a, and the circumbinary disk left behind, as the binary spirals inward fast – e.g. e.g. Pringle, 1991; Armitage+2002, Milosavljevic+2005.

- Merger simulations in full numerical relativity hint at interesting dynamics, but too short … e.g.Bode+2010; Farris+2010, Farris+2011, Giacomazzo+2012; Gold+ 2013; Kelly+2017.

- Modern 2D hydrodynamics hints a lot of accretion! – D'Orazio+ 2013; Farris +2014; Ryan+2016, Tang+2018.

- Modern 3D GR-MHD now have reversed completely the picture!– e.g. Shi+12, Noble+12, Bowen+18,19.

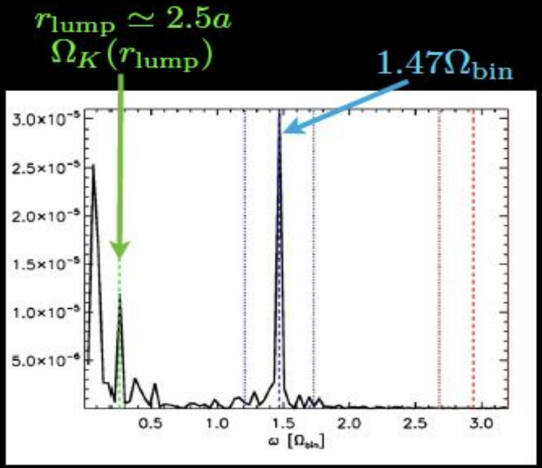


Shi+ 2012

Binary torque "dam" does not hold, and accretion continues until approach to merge!

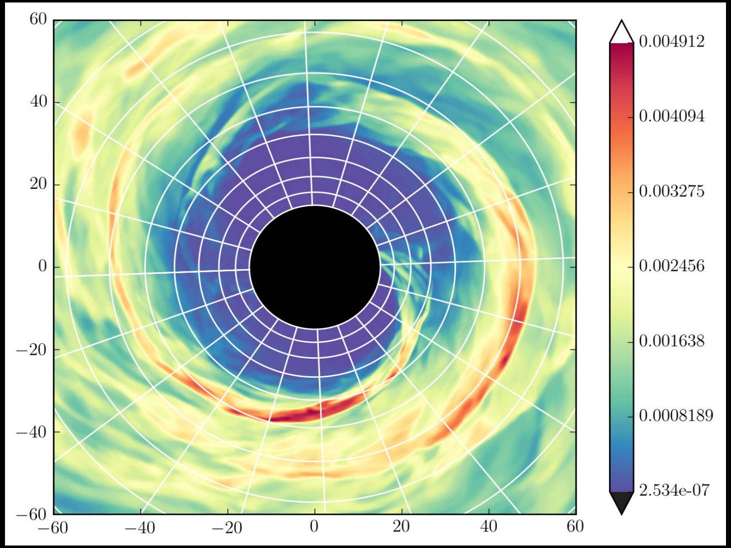
Circumbinary Disk Dynamics

We found dense **accretion streams** to the BHs, and **overdensity** ("lump") in the circumbinary disk with characteristic EM signal periodicity $\Omega_{beat} = \Omega_{bin} - \Omega_{lump}$ – Shi+12, Noble+2012



Noble, Mundim, Krolik, Campanelli + ApJ 2012

Long term MHD simulations (equal-mass)
(BHs not on the grid, initial BH sep.=$20r_g$)



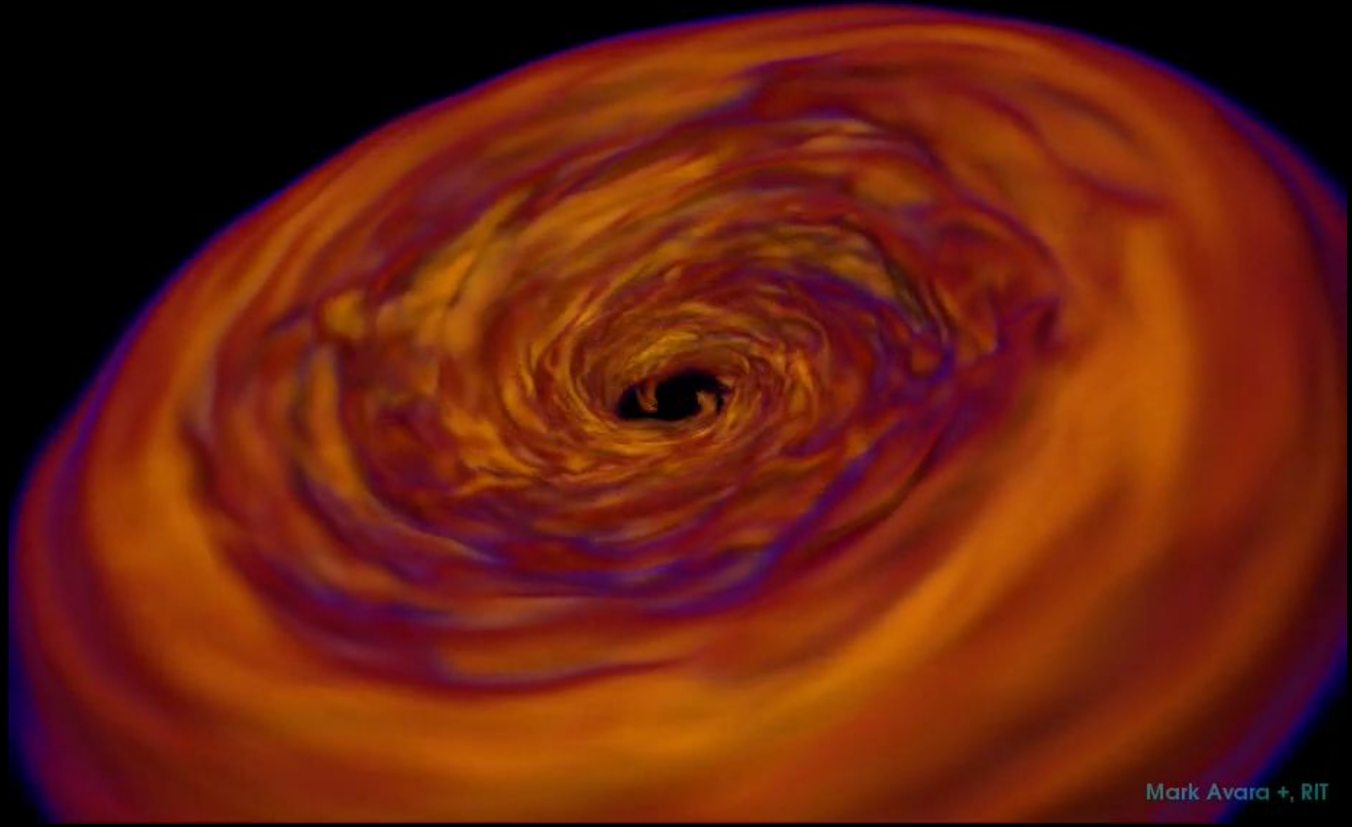
Noble +2012

This qualitative picture holds for nearly equal mass BHs (q >= 1/5), and is independent of disk size or magnetization – Noble+, in prep 2019

Do not see a lump for ~1:10 mass ratio!

Circumbinary Disk Dynamics

Long term MHD simulations
of a tilted cicumbinary disk (~12 deg)
initial BH sep=43r_g, final BH sep=8r_g
(BHs not on the grid) – Avara+2019 in prep



Long-term GR-MHD simulations

Gas evolution through conservation of mass, energy and momentum, and Maxwell's equations, on dynamical binary BH spacetime:

$$\frac{\partial}{\partial t}\mathbf{q}(\mathbf{P}) + \frac{\partial}{\partial x^i}\mathbf{F}^i(\mathbf{P}) = \mathbf{S}(\mathbf{P})$$

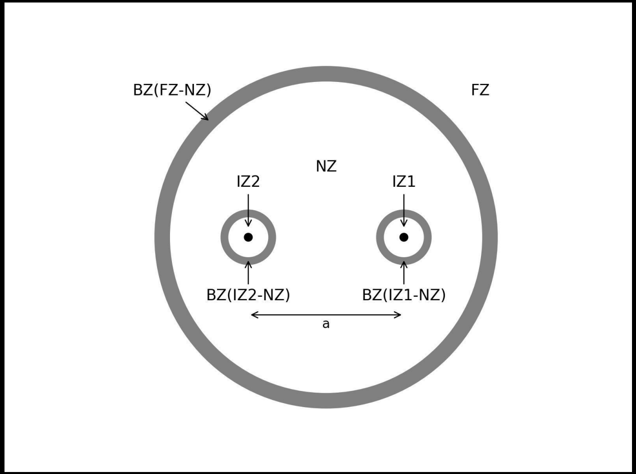
$$\frac{\partial}{\partial t}\sqrt{-g}\begin{bmatrix}\rho u^t \\ T^t{}_t + \rho u^t \\ T^t{}_j \\ B^k\end{bmatrix} + \frac{\partial}{\partial x^i}\sqrt{-g}\begin{bmatrix}\rho u^i \\ T^i{}_t + \rho u^i \\ T^i{}_j \\ \left(b^i u^k - b^k u^i\right)\end{bmatrix} = \sqrt{-g}\begin{bmatrix}0 \\ T^\kappa{}_\lambda \Gamma^\lambda{}_{t\kappa} - \mathcal{F}_t \\ T^\kappa{}_\lambda \Gamma^\lambda{}_{j\kappa} - \mathcal{F}_j \\ 0\end{bmatrix}$$

$$T_{\mu\nu} = \left(\rho + u + p + 2p_m\right) u_\mu u_\nu + \left(p + p_m\right) g_{\mu\nu} - b_\mu b_\nu$$

Mass Density (ρ); Internal Energy Density (u); Gas Pressure (p); Fluid's 4-velocity (u_μ); Magnetic Pressure (p_m); Magnetic 4-vector (b_μ); Radiative Energy & Momentum Loss ($\mathcal{F}$)

- Use a well-tested, **flux-conservative**, generally covariant, GRMHD code for BH accretion disks: Harm3D – Gammie, McKinney & Toth 2003, Noble+2006
  - Ideal gas (polytropic +piecewise EOS)
  - **Isentropic cooling** (to target S_0) to keep H/r ~constant

- Code adapted to handle dynamical gravity in the relativistic GW inspiral regime – Noble+2012, Mundim+2014, Ireland+2014
  - Binary BH spacetime valid for any mass ratio and BH spins at a given initial separation.
  - BHs inspiral via the Post-Newtonian equations of motion.



Computational Strategies:

Evolve accreting inspiraling BH binaries while **resolving the MRI and MHD dynamics** at the scale of the event horizons:

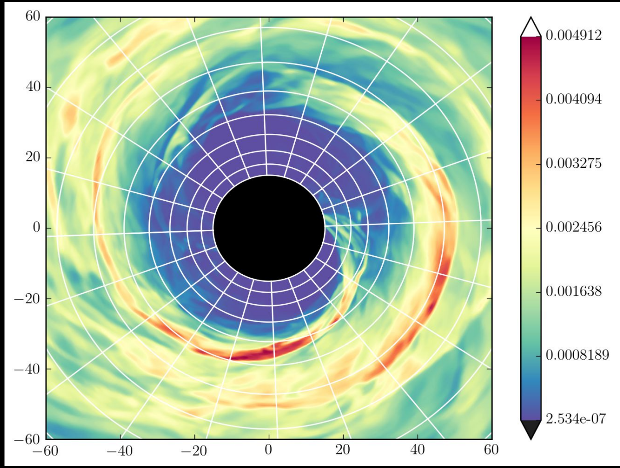
1. Perform a long-term GRMHD simulation of a thin, radiatively efficient, circumbinary accretion disk to its "quasi-steady" state:

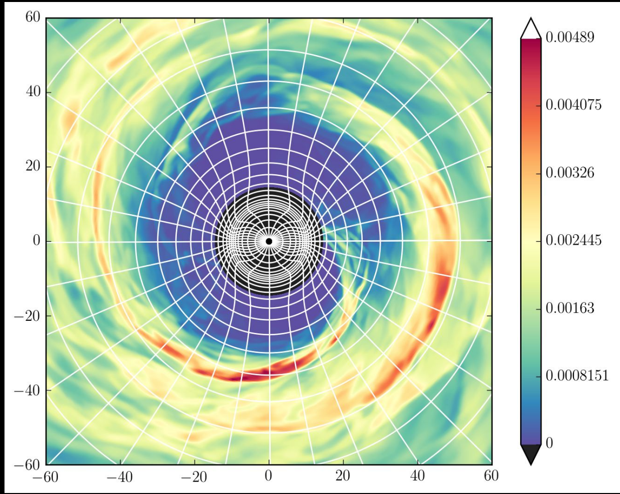
    - Use spherical polar, horizon penetrating, coords for proper angular momentum transport in the gas;
    - Remove the BHs from the grid for efficiency at this stage;

    This allow us to follow the circumbinary disk MHD dynamics for hundreds of orbits as the binary approach merger!

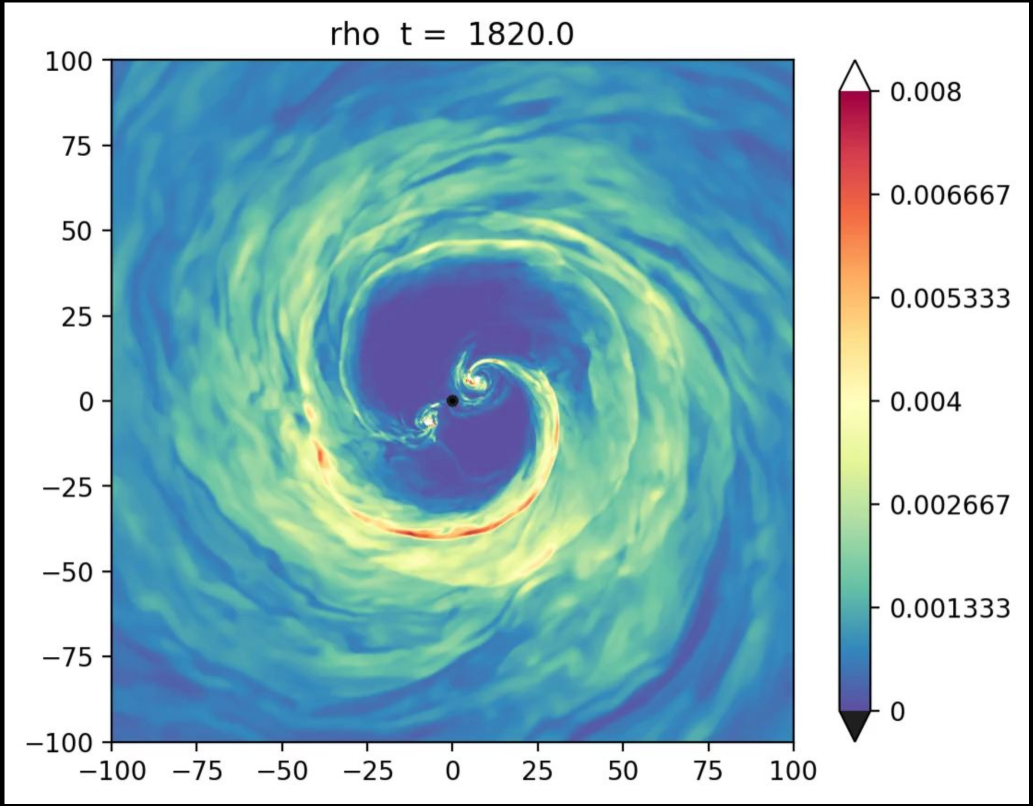
2. At "equilibration", interpolate the computational domain into a new grid designed to resolve the physics near each BHs:

    - Novel methods tailored for accuracy and efficiency e.g. dynamics warped grid – Zilhao+2014;
    - Now, augmented efficiency with a new multipatch code – Avara+2019



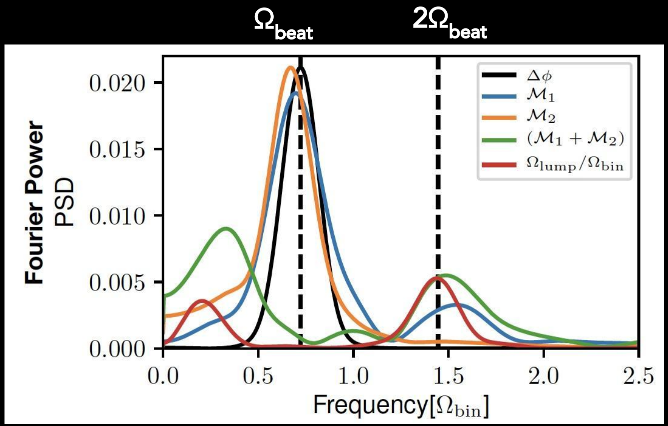


Dynamics in the Central Region



We discovered new dynamical interactions between the mini-disks and circumbinary disk – Bowen+ ApJL 2018, Bowen+ ApJ 2019

- Accreting streams fall in the cavity and shock against the individual BH mini-disks.
- Mini-disks deplete and refill the disks periodically at time scale close to one orbital period.



Bowen, Mewes, Campanelli, Noble, Krolik, ApJL 2019

Bowen, Mewes, Noble, Avara, Campanelli, Krolik, ApJ 2019

General Relativistic Radiative Transfer:

- **Bothros** - General relativistic ray-tracer for transporting radiation emitted from 3D GR-MHD simulation snapshots – Noble+2009
  - Radiative transfer integrated back into the geodesics
  - Local cooling rate = local bolometric emissivity
- **Thermal Photosphere**:
  Photons starting at photosphere start as black-body

$$\frac{\partial I}{\partial \lambda} = j - \alpha I \qquad I_\nu = B_\nu(\nu, T_{\rm eff}) = \frac{2h\nu^3}{c^2}\frac{1}{e^{\frac{h\nu}{kT_{\rm eff}}} - 1}.$$

Opacity: grey Thomson opacity for electron scattering

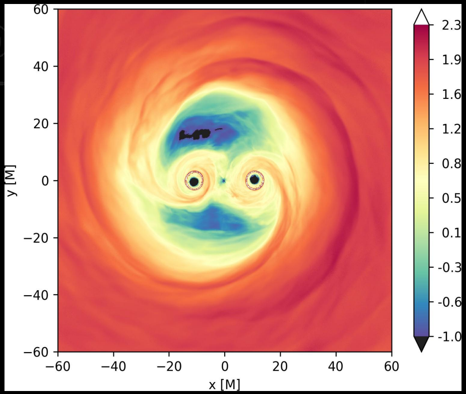
- **Above photosphere, corona emission** modeled as non-thermal (Compton scattering) component with temperature 100 keV:

$$j_\nu \propto \mathcal{W}_\nu = \left(\frac{h\nu}{\Theta}\right)^{-1/2} e^{-\frac{h\nu}{\Theta}} \qquad \Theta = kT/m_e c^2 = 0.2$$

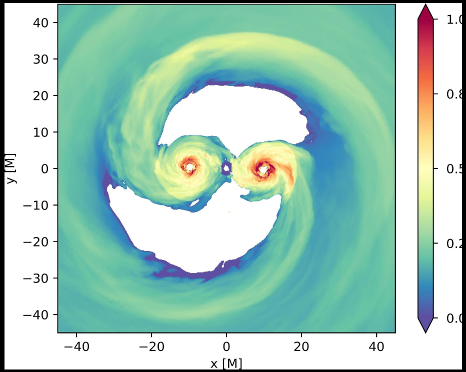
Trakhtenbrot++2017, Krolik 1999, Roedig++2014

- Emissivity ignored in low-density regions in which scattering processes are important (and unavailable to us for now);

- Explore opt. thin and thick cases: $\dot{m} = 8 \times 10^{-4}$, $\dot{m} = 0.5$



Log10 Optical Depth
Grey Thomson Opacity

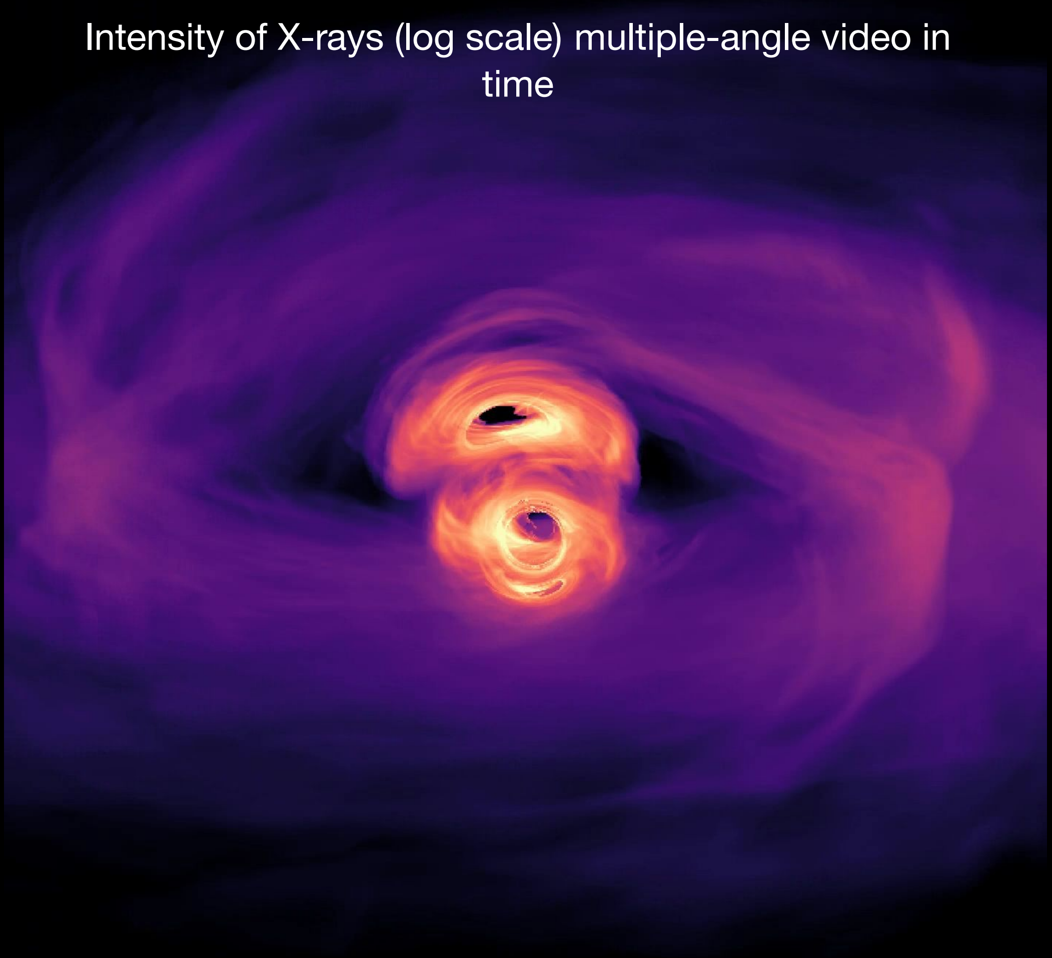


Map of Photosphere's Location & Temperature

$\mathrm{Log}_{10}(T_{\rm eff}/T_0)$, $T_0 = 5 \times 10^5$K

Optically Thick Case

Intensity of X-rays (log scale) multiple-angle video in time



Credits: S. Noble (NASA) based on Bowen+2018

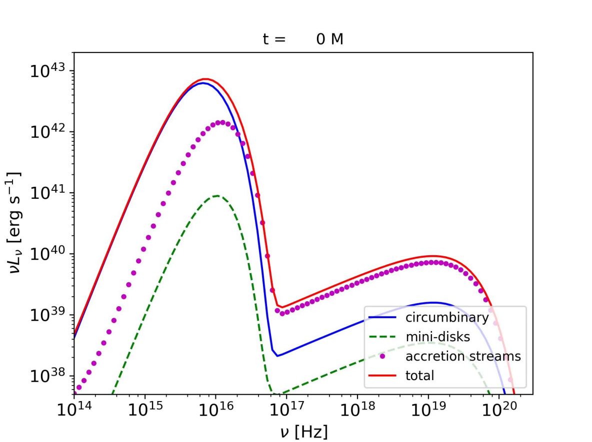
Calculations of Distinct Light Signatures

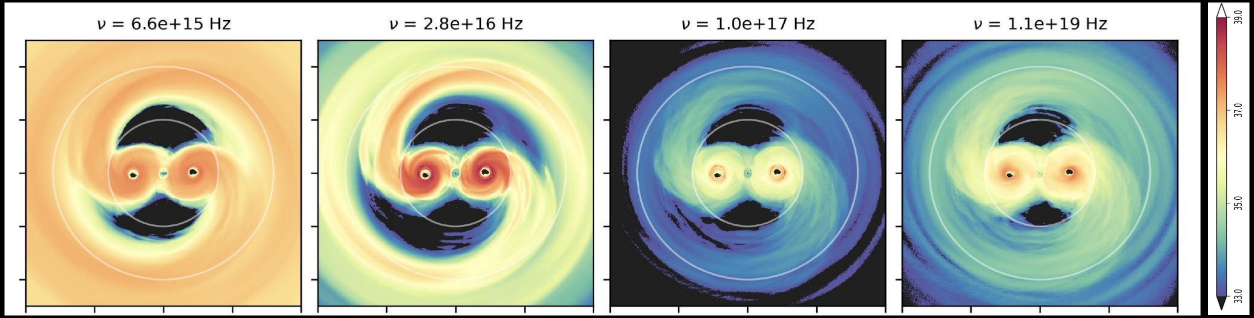
The first predicted time varying spectrum from accreting binary black holes in the inspiral regime – D'Ascoli+2018

Key distinctions from single BH (AGN) systems:

- Brighter X-ray emission relative to UV/EUV.
- Variable and broadened thermal UV/EUV peak.
- "Notch" between thermal peaks of mini-disks and circumbinary disk – e.g. e.g. Roedig+2014 – will likely be more visible at larger separations and for spinning black holes.

Face-on View,
Optically Thick Case
$M_{BH}=10^6 M_\odot$





Circumbinary dominated UV

Mini-disk dominated soft X-rays

X-rays near the boundary between thermal and corona dominance

Mini-disk corona dominated hard X-rays

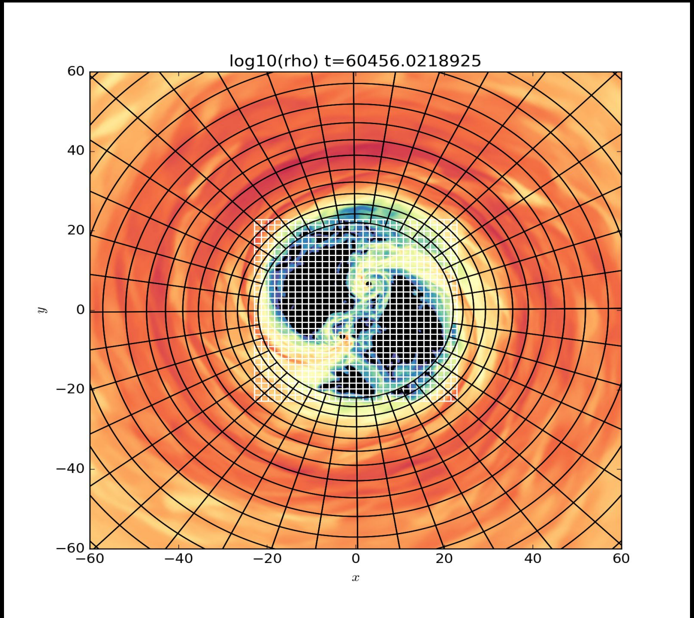
The systems will likely be too distant to be spatially resolved, so we need to understand their spectrum and how it varies in time.

D'Ascoli, Noble, Bowen, Campanelli, Krolik, ApJL 2019

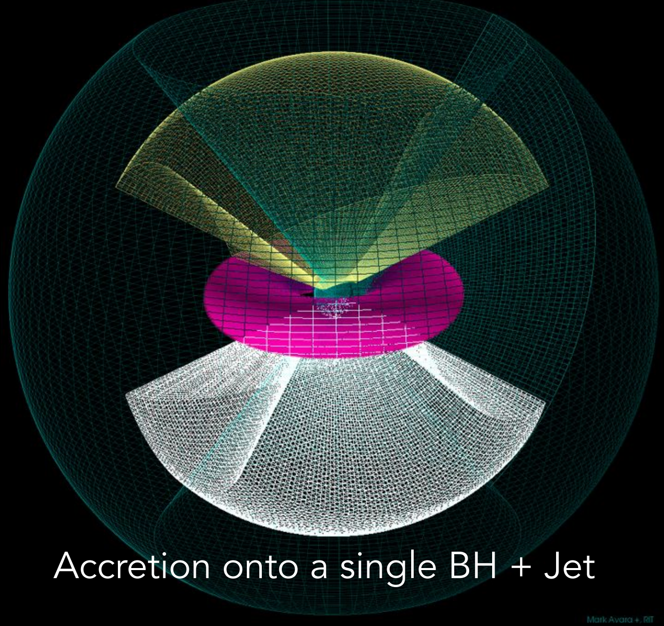
New software infrastructure for problems of discrepant physical, temporal, scales and multiple geometries

How do we efficiently simulate 10^7-10^8 cells for 10^6-10^7 steps?

- PatchworkHD– Shiokawa+ 2018; **PatchworkMHD** – Avara+ 2019 in prep

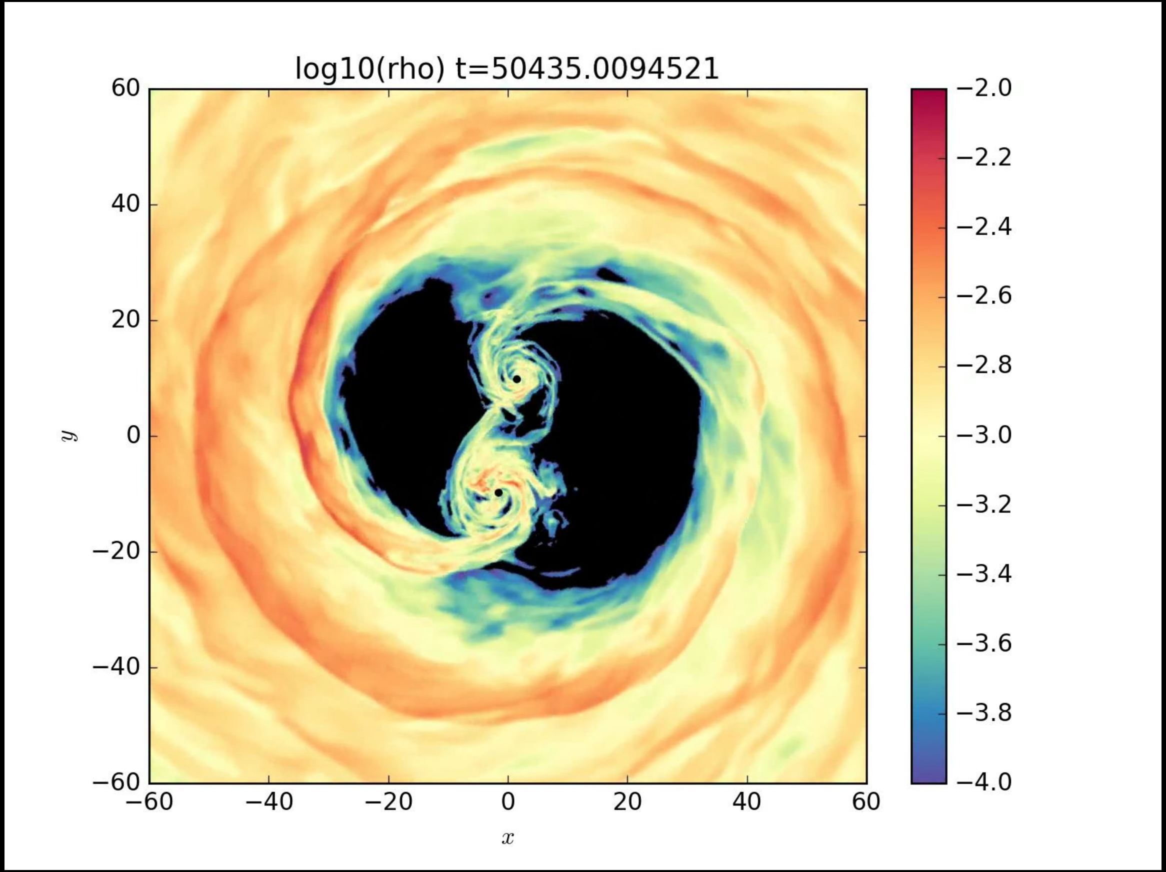


Accretion onto binary BHs



Accretion onto a single BH + Jet

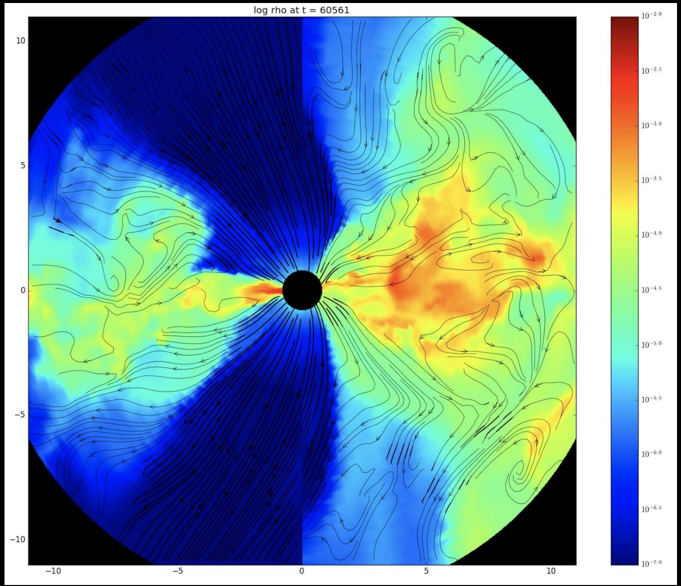
First Long-term Global MHD Dynamics



Avara+2019, in prep

Ongoing PatchworkMHD Simulation –

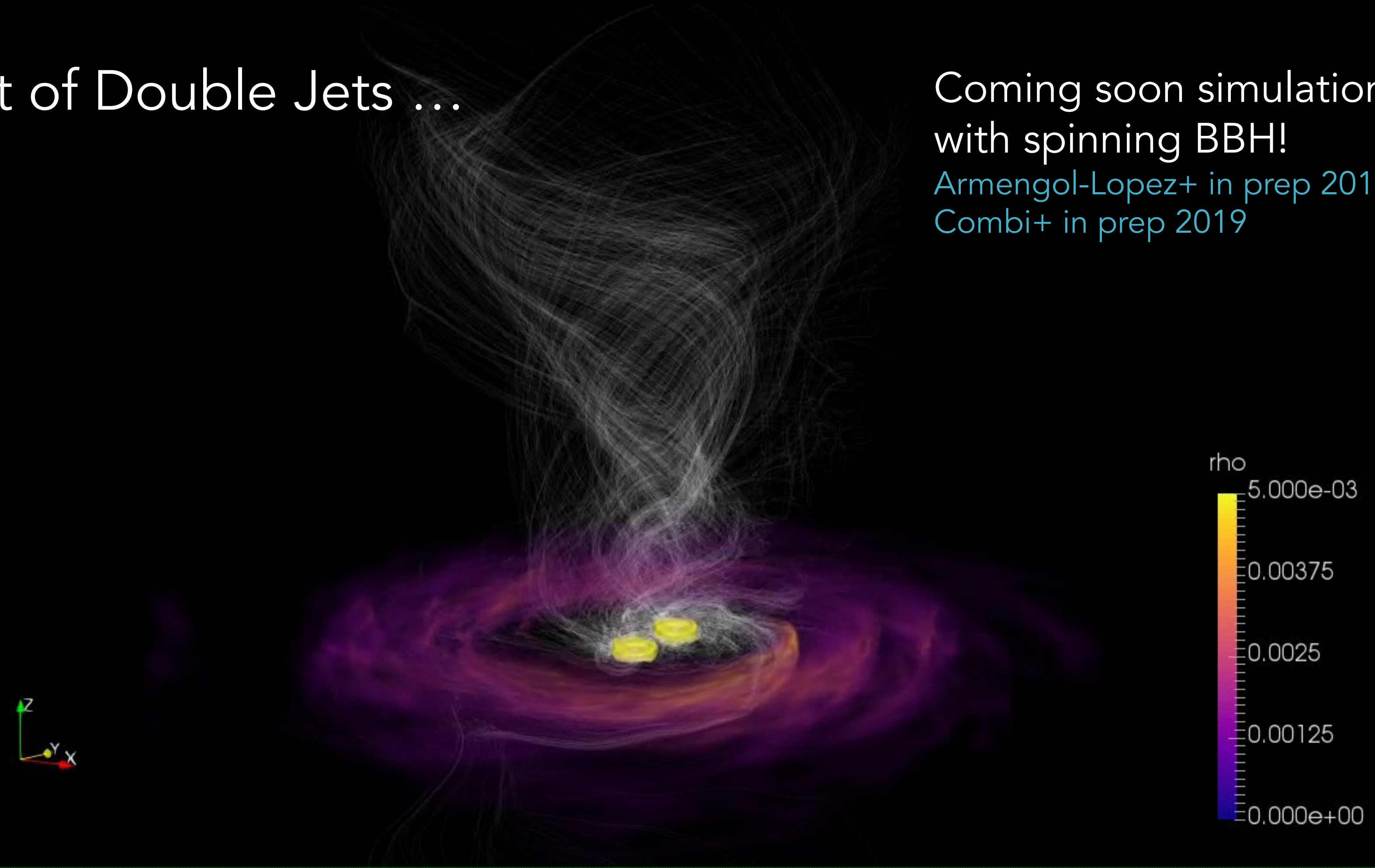
- First physical parameter studies of these systems in 3D GRMHD
- Now **30 times our prior efficiency**
- Sufficient time series data to calculate **light curve** (being analyzed now)



Hint of Double Jets …

Coming soon simulations with spinning BBH!

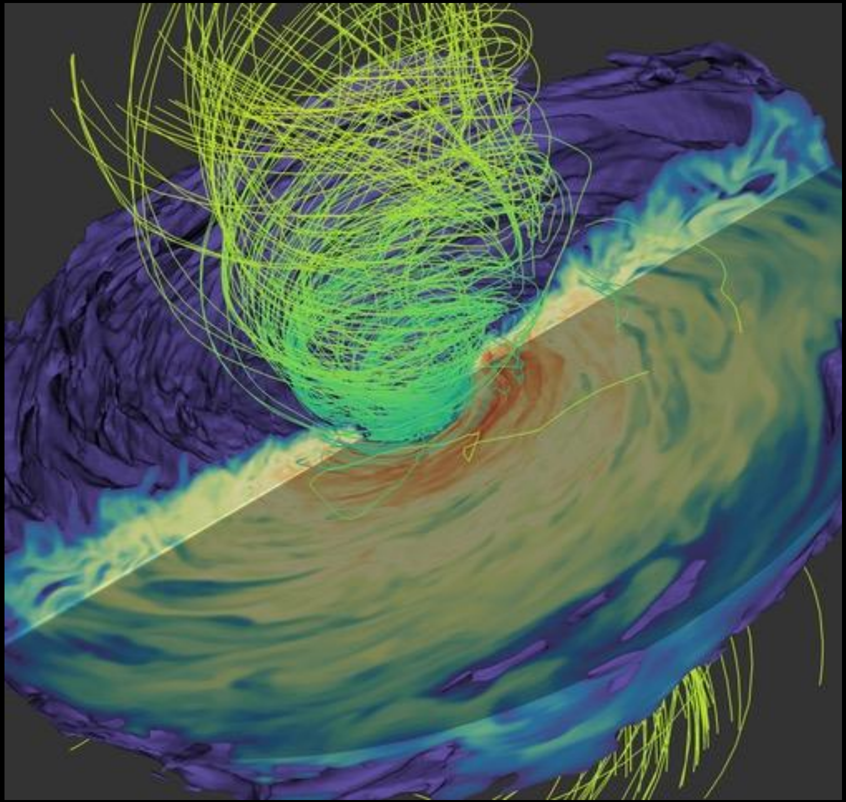
Armengol-Lopez+ in prep 2019
Combi+ in prep 2019



Bowen++, ApJL, 2018 ; Bowen++, ApJ, 2019

Summary

- Binary black hole mergers, in particular supermassive mergers are ideal multi-messenger sources!
- Simulations of galaxy evolutions predicts a non-negligible fraction of these sources within the PTA (and LISA) range.
- Accurate 3d GRMHD models are now accurate and long enough to predict distintive EM signals!
- Black holes are “hot”, but there is still a lot that we don’t know about them. There might be surprises awaiting for us!



Credits: Mewes+, RIT 2019