The splashback radius as a physical halo boundary

Benedikt Diemer

Einstein Fellow Harvard-Smithsonian CfA / University of Maryland

(in collaboration with Andrey Kravtsov, Surhud More, Philip Mansfield, Spencer Scott, Alexie Leauthaud, and many others)

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R_{500c}

Visualization code: Phil Mansfield





Visualization code: Phil Mansfield



 $\mathsf{R}_{\mathsf{vir}}$

Visualization code: Phil Mansfield



R_{200m}

Visualization code: Phil Mansfield





"Ejected" satellites



Balogh et al. 2000 • Mamon et al. 2004 • Gill et al. 2005 • Wetzel et al. 2014

Pseudo-evolution



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Diemer et al. 2013ab • Cuesta et al. 2008 • Zemp 2014 • More et al. 2015 • Wetzel & Nagai 2015



Fillmore & Goldreich 1984 • Bertschinger 1985 • Lu et al. 2006 • Diemand & Kuhlen 2008 Vogelsberger et al. 2011 • Lithwick & Dalal 2011 • Adhikari et al. 2014 • Shi 2016



Low accretion rate

High accretion rate

Diemer & Kravtsov 2014 • Adhikari, Dalal & Clampitt 2014 • More, Diemer & Kravtsov 2015

Large halos (M>10¹⁵)



Splashback in the real Universe



Chang et al. 2017 • Diemer & Kravtsov 2014 • Navarro et al. 1997

Splashback in observations and simulations



Simulations: More et al. 2015 • Diemer et al. 2017

Observations: More et al. 2016 • Baxter et al. 2017 • Chang et al. 2017 • Shin et al. 2019 • Zuercher & More 2019 See also: Tully 2015 • Zu et al. 2017 • Busch & White 2017 • Patej & Loeb 2016 • Umetsu & Diemer 2017

eROSITA X-ray selection



TA

Cosmology with the splashback boundary?



Splashback and SIDM



Euclid











Spencer Scott

Philip Mansfield

Scott, Mansfield, Diemer 2019 (in prep.) • Mansfield, Kravtsov & Diemer 2017 • Penna & Dines 2007

Conclusions

- The structure of CDM halos is not a solved problem, particularly wrt their dynamics
- The **splashback radius** provides a physical halo boundary
- Observations of the halo **outskirts** can tell us about the physics of dark matter and gravity